THE SOLAR CYCLE AND THE TACHOCLINE: THEORIES AND OBSERVATIONS

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ABSTRACT

We present some of the theories of what Petrovay (2001) called the DOT (Dynamo-Overshoot-Tachocline) layer emphasizing some predictions that may be tested by helioseismology. We first recall some constraint on these theories obtained by the previous latitudinal-average studies of the tachocline. Then, we investigate the latitudinal and temporal variations of the tachocline parameters by using our recent analysis of the 6 years of helioseismic observations with the LOWL instrument operating in Mano Loa, Hawaï since 1994. Within the observational uncertainties, our preliminary work yields no compelling evidence of variations or general trend of the tachocline parameters during the ascending phase of the current solar cycle (1994-2000).

Key words: Sun; Tachocline; Overshoot; Solar Cycle.

1. INTRODUCTION

The solar tachocline, so called after Spiegel & Zahn (1992), is defined as the layer at the base of the convection zone (hereafter, CZ) where important radial shear occurs. It is the transition zone where the angular velocity changes from its latitude-dependent value in the CZ to its constant and intermediate value in the radiative interior. We can emphasize three main reasons why this layer is of particular interest.

(i) Shear turbulence and/or meridional circulation inside the tachocline may provide a mechanism for mixing material between the CZ and the radiative interior (e.g. Brun et al. 1999; Schatzman et al. 2000) which is needed to understand, for instance, the burning of lithium and the depletion of helium and to reach a better agreement between solar models and helioseismic observations (e.g. Richard et al. 1996; Brun et al. 1999).

(ii) The tachocline may be the seat for angular momentum transport processes that could lead to the observed rigid rotation rate of the radiative interior. Hydrodynamical transport by unstable shear flow (e.g. Chaboyer et al. 1995) and transport by internal gravity waves in the tachocline (Kumar et al. 1999; Kim & MacGregor 2000) have been studied but they have not been found efficient enough to lead to an uniform internal rotation at the present age of the Sun. This probably requires also the presence of an internal magnetic field (Mestel & Weiss 1987; Charbonneau & MacGregor 1993; Gough & McIntyre 1998).

(iii) Finally, the tachocline is the best location for an oscillatory solar dynamo which is generally believed to be responsible for the solar magnetic cycle for the following reasons: (1) Its radial and latitudinal differential rotation has the ability to produce a toroidal field by shearing a pre-existing poloidal field. (2) The α-effect, the essential mechanism for producing poloidal field from toroidal field, is usually located in the CZ but the tachocline can also produce a strong α-effect by magnetic buoyancy instability (Ferriz-Mas et al. 1994) and/or by the unstable shallow-water modes (Dikpati & Gilman 2000a). (3) Because the tachocline (or part of it) may also be located in the slight sub-adiabatic overshoot layer, the toroidal fields can be stored for an extended period of time, and therefore can be amplified and acted on by the α-effect before they escape to the surface through buoyant rise or be disrupted completely by convective shredding.

We will first (Sect. 2) give some terminology about the different layers and their properties at the base of the CZ (also to be defined). Then, in Sect. 3, we present some models of the tachocline and their predictions that may be tested by helioseismology. Previous results obtained by forward modeling on the latitude-average properties of the tachocline are summarized in Sect 4. Then we use a new dataset from the LOWL instrument to investigate the latitudinal and temporal variations of the tachocline parameters during the ascending phase of the current solar cycle (Sect. 5).

2. LAYERS AT THE BASE OF THE CONVECTION ZONE: TERMINOLOGY, DEFINITIONS AND OBSERVATIONS

As mentioned above, there is, beside the tachocline, another layer defined at the interface between the superadiabatic CZ and the radiative interior: the overshoot layer. Different models exist for this layer. The first

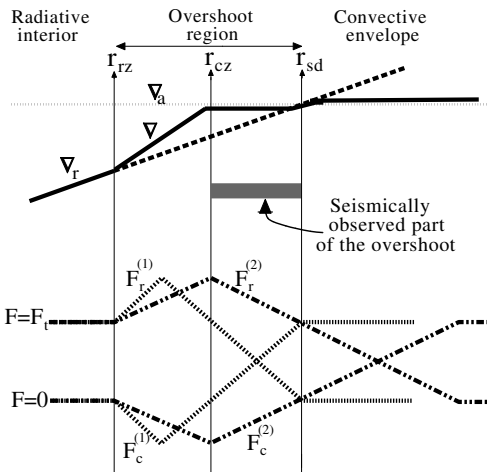

Figure 1. Schematic representation of the overshoot layer between the radiative zone boundary r_{rz} and the Schwarzschild boundary r_{sd} (i.e. where the temperature gradient equals both its adiabatic and radiative value $\nabla = \nabla_a = \nabla_r$). r_{cz} represents the seismically observed base of the CZ. Two different models of the convective (F_c) and radiative (F_r) flux variations are drawn for comparison. The total flux $F_t = F_c + F_r$ remains constant. The model (1) of Canuto (1997) allows a positive convective flux in the slightly sub-adiabatic and seismically observed part of the overshoot in contradiction with the Schwarzschild criterion while the model (2) of Zahn (1991) is such that $F_c \leq 0$ in all the overshoot layer.

particular radius, r_{sd}, corresponds to the Schwarzschild boundary which defines the beginning of the convectively unstable zone where the temperature gradient equal its adiabatic value $\nabla = \nabla_{ad}$. A second particular radius is where the temperature gradient is taking its radiative value ($\nabla = \nabla_r$) and the total flux is equal to the radiative flux ($F_t = F_r$). This radius, r_{rz}, defines the radiative zone boundary. The region between r_{sd} and r_{rz} is what we here call the overshoot region (see Fig. 1). The 'base of the convection zone' r_{cz}, measured by helioseismology, corresponds to the location where the temperature gradient changes abruptly. If we allow a slightly sub-adiabatic zone below r_{sd} then helioseismology cannot distinguish between the slightly super-adiabatic CZ and the slightly sub-adiabatic overshoot layer (Christensen-Dalsgaard et al. 1995; Gilman 2000) and $r_{rz} < r_{cz} < r_{sd}$. Calibrations made by Basu (1997) lead to a value of $r_{cz} = 0.713 \pm 0.001 R_\odot$. Helioseismology can also place an upper limit for the slightly sub-adiabatic overshoot zone $[r_{cz}, r_{sd}]$. This upper limit is 0.05 pressure scale heights, about 2800Km or $\sim 0.004 R_\odot$ (Basu & Antia 1997). This is therefore a very thin layer. Another question concerning this layer is about the sign of the convective flux inside. If we follow the Schwarzschild criterion, we should have $F_c \leq 0$ as in Zahn (1991)'s model because this layer is (slightly) sub-adiabatic, but Ferriz-Mas (1996) and Canuto (1997) pointed out that, with a nonlocal mixing-length treatment of the CZ, this is not necessarily valid in the overshoot zone and we may well have $F_c > 0$ and $\nabla < \nabla_{ad}$ in spite of the Schwarzschild criterion. In that case the downward convective energy transport ($F_c < 0$) occurs only in the other (lower) part of the overshoot zone $[r_{rz}, r_{cz}]$, called the overshoot layer proper (Canuto 1997; Monteiro et al. 1998). In the same way, the zone with positive convective flux is sometimes called CZ proper (e.g. Ferriz-Mas 1996). These differences in the overshoot models may have some importance as it has been shown that flux tubes with equipartition field strength ($\sim 10^4$ G) can be stored in all the overshoot layer while the field strength of 10^5 gauss required in order to explain the observed active regions at the surface (e.g. Schüssler et al. 1994), can only be stored where the convective flux is negative (Ferriz-Mas 1996).

3. THEORY: SOME MODELS OF THE TACHOCLINE AND THEIR PREDICTIONS

3.1. Purely hydrodynamic models

The first theory of the tachocline was developed by Spiegel & Zahn (1992). They mainly address the problem of the thickness of the tachocline, which we summarized as follows. The differential rotation $\Omega_1(\theta)$ at the top of the radiative interior induces a latitudinal temperature gradient and therefore a meridional circulation that would, without any stress acting on it, spread towards the interior in a thermal time-scale leading to a non rigid rotation rate in the interior and a very thick tachocline for the present Sun. We therefore need to invoke some stresses acting in the tachocline in order to prevent its progression towards the interior. Using a purely (i.e. non-magnetized) hydrodynamic model where the tachocline is treated as a boundary layer, Spiegel & Zahn (1992) suggest that a strong enough anisotropic turbulence could lead to a thin tachocline. They thus obtained a relation between the thickness of the tachocline w and the horizontal turbulent viscosity coefficient ν_h assumed to be much higher than its vertical counterpart.

$$\nu_h = 8.34\ 10^6 \left(\frac{\Omega_0}{N}\right)^2 \left(\frac{r_{cz}}{w}\right)^4 \text{ cm}^2\ \text{s}^{-1} \qquad (1)$$

Nevertheless there remain two major questions with this approach. First, the latitudinal rotation profile inside the tachocline is likely to be linearly stable to 2D disturbances according to Rayleigh's criterion (Charbonneau et al. 1999b) and so it is probably not the process leading to the anisotropic turbulence assumed in the model. A solution to this problem may be found in a recent work of Dikpati & Gilman (2000a) showing that, if we allow radial deformations of the layer and use the so-called shallow water model as a first approximation of the full 3D problem, then the tachocline is found unstable in the slightly sub-adiabatic overshoot zone even in the purely hydrodynamic case (Fig. 2). The second difficulty with this hydrodynamic model is that, even if it leads to a latitudinal independent rotation at the base of the tachocline, it cannot explain the fact that there is also no significant radial gradient in the interior.

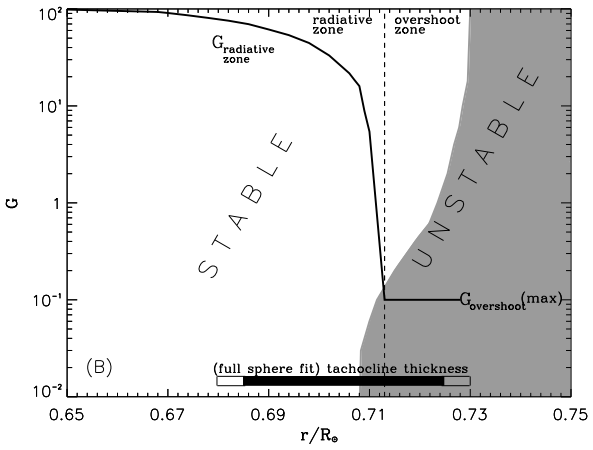

Figure 2. Stability diagram from the shallow water model (Dikpati & Gilman 2000a). The variation of $G \simeq 10^3|\nabla - \nabla_{ad}|$ with radius is given by the solid line for schematic solar model. The stability zones and the tachocline thickness and location have been estimated from observations (Charbonneau et al. 1999a,b). The overshoot layer has arbitrarily been assigned a thickness of $0.015R_\odot$ and it is slightly sub-adiabatic with $|\nabla - \nabla_{ad}| \leq 10^{-4}$.

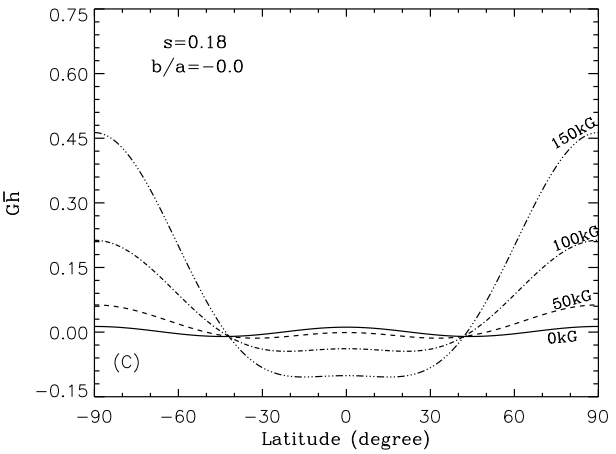

Figure 3. Shape of the top surface of the tachocline with a broad nonuniform toroidal magnetic field of different strength. G is the same quantity as in Fig. 2 and h is related to the width of the tachocline by $w \sim 0.035R_\odot(1+h)$ Dikpati & Gilman (2000b).

3.2. Magnetized models

Because it is very likely that the tachocline is magnetized, another mechanism that can lead to anisotropic turbulence in the tachocline is magnetic instability. It has effectively been shown (Gilman & Fox 1997, 1999; Dikpati & Gilman 1999) that there exists a joint instability between latitudinal differential rotation and toroidal magnetic fields. But two different kind of magnetic fields are usually invoked in the solar interior, namely primordial fields and dynamo generated fields. As we shall see, in both cases theories relate their properties to the shape of the tachocline.

(i) Dynamo generated magnetic fields. In virtually all dynamo theories (overshoot layer, interface or Babcock-Leighton dynamos, see the review of Petrovay (2001)), the shape and especially the thickness of the tachocline and the overshoot layer are key issues since they determine the strength of the magnetic field that can be stored and the process by which it is transformed. Using an MHD version of their shallow water model (Dikpati & Gilman 2000b) demonstrate that the presence of a magnetic field in the tachocline would induce a prolateness of the tachocline otherwise oblate. This can also be tested by helioseismology by measuring the central position of the radial shear at different latitudes. Such measurement of the prolateness would allow us to estimate the strength of the magnetic field depending on its geometry and localization inside the tachocline. The shape of the tachocline as obtained from this model is shown as a function of the toroidal field strength on Fig 3.

(ii) Primordial magnetic fields. Independently of any magnetic field that may be generated in the CZ, Gough & McIntyre (1998) argue that we must also have a magnetic field in the radiative interior in order to explain the uniform (with no radial shear) rotation rate observed there. This large scale poloidal magnetic field would also confine the shear in a thin layer as a result of a balance between upward diffusion of the magnetic field and downward advection by the thermally driven tachocline circulation. They obtain the following relation between the thickness of the tachocline and the strength of the internal magnetic field:

$$|B_0| = 2.3\ 10^3 \left(\frac{\Omega_0}{N}\right)^7 \left(\frac{r_{cz}}{w}\right)^9 \text{ (Gauss).} \qquad (2)$$

Such internal field is also required in the magnetic models of Rudiger & Kitchatinov (1997) and MacGregor & Charbonneau (1999) where the tachocline is identified with an MHD boundary layer located in the radiative interior. Assuming no magnetic coupling at the core-envelope interface and that advective effects (as modeled by Spiegel & Zahn (1992)) dominate the viscous effects to prevent the inward spreading of the layer, MacGregor & Charbonneau (1999) obtain:

$$|B_0| \sim 3\ 10^{-8} \left(\frac{r_c}{w}\right)^3 \text{ (Gauss).} \qquad (3)$$

This work also suggests that the primordial magnetic field is very likely to be decoupled from the tachocline i.e. entirely contained within the radiative interior. The interaction between a primordial time-independent magnetic field that would extend into the CZ and the time-dependent dynamo-generated magnetic field had nevertheless been investigated by Boruta (1996) suggesting that this interaction would lead to an inversion and an amplification of the primordial field by an amount that depends again on the thickness of the tachocline. On the other hand, if the dynamo magnetic field is generated in the CZ, it has been shown to have only a weak influence on the dynamics of both the tachocline and the radiative interior (Hujeirat & Yorke 1998; Garaud 1999).

4. HELIOSEISMOLOGY: FORWARD MODELING AND THE LATITUDE-AVERAGE TACHOCLINE

Solar oscillation modes are identified by three integers: the spherical harmonics degree l and azimuthal order m, and the radial order n. The frequency splitting $\delta\nu_{nlm}$ induced by the rotation $\Omega(r,\mu)$ can be calculated by:

$$\delta\nu_{nlm} = \int_0^1 \int_0^1 K_{nlm}(r,\mu)\ \Omega(r,\mu) dr d\mu \qquad (4)$$

where, r is the fractional solar radius, $\mu = \cos(\theta)$ and K_{nlm} are kernels calculated from a standard solar model.

Following Kosovichev (1996) we parameterize the rotation rate between 0.4 and 0.8 solar radius by an error function of the form:

$$\Omega(r,\theta) = \Omega_0 + \frac{\Omega_1(\theta) - \Omega_0}{2} \left(1 + erf\left(\frac{r - r_c(\theta)}{0.5\ w(\theta)}\right)\right) \quad (5)$$

where Ω_0 represents the rigid rotation rate in the radiative interior, $\Omega_1(\theta)$ the rotation rate at the top of the tachocline, and $r_c(\theta), w(\theta)$ the central position and the width of the tachocline at the colatitude θ.

Several authors have inferred the latitude-average properties of the tachocline (i.e. r_c and w independent of θ) by looking for the parameters that give the best fit to the helioseismic observations, either by calibration (Kosovichev 1996; Basu 1997) or by genetic forward modeling (Charbonneau et al. 1999a). A weighted average of these results leads to a position $r_c = 0.704 \pm 0.002 R_\odot$ and a width $w = 0.052 \pm 0.010 R_\odot$. Therefore, it seems that, in a latitude-average, the tachocline is centered around $0.01 R_\odot$ below r_{cz} and its upper third part is located in the nearly adiabatically stratified zone above r_{cz}. Nevertheless this results must be taken with caution because the decreasing radial resolution with latitude tends naturally to show a thicker tachocline at high latitudes thereby influencing the average thickness and position. The parameters obtained this way should therefore be regarded as upper limits for the tachocline properties. The latitude-average width is however probably the best estimate to use in the formulae obtained from the various theories of Sect. 3 which assume also a spherical symmetry.

In the hydrodynamic theory, Eq. 1 with a with of $0.05 R_\odot$ leads to an horizontal turbulent viscosity coefficient of about 3 10^6cm^2 s^{-1}, several order of magnitude higher than the microscopic value. In the MHD theories the same value of the width leads to a primordial magnetic field strength of $|B_0| \sim 10^{-4}$Gauss for both Eq. 2 and 3. This suggests that even a weak magnetic field is enough to keep the radiative interior rotating rigidly and to confine the radial shear to a thin layer compatible with observations. By analyzing the splitting coefficients from GONG observations, Basu (1997) set an upper limit of 0.3MG for a field located at the base of the CZ. From Eq. 2 this would correspond to a width of $0.0045 R_\odot$ which cannot be excluded from observations. By shearing a poloidal field of 10^{-4}Gauss the theory of MacGregor & Charbonneau (1999) predicts the generation of a toroidal field of ~ 0.1MG which is also compatible with the 0.3MG limit.

5. SEARCH FOR VARIATIONS WITH TIME AND LATITUDE USING INVERSION TECHNIQUES

5.1. Motivation

Theories predict variation with latitude of the shape of the tachocline. We have seen for example that, in the shallow-water model, the presence of a magnetic field would induce a prolate tachocline otherwise oblate. Both the magnetic model of Rudiger & Kitchatinov (1997) and the 2D hydrodynamic model of Spiegel & Zahn (1992) predict a thicker tachocline at the pole than at the equator (see also Elliott 1997). In the model of Gough & McIntyre (1998) the tachocline is thicker where the horizontal component of the magnetic field in the radiative interior is small. The poles should be therefore thicker but that may also be true for other latitudes. This means that the shape of the tachocline could also be a diagnostic for the geometry of the magnetic field. Some attempts have been made to infer latitudinal variations in the tachocline profile from observations (Antia et al. 1998; Charbonneau et al. 1999a), but no compelling evidence has been found yet. Even if, as shown in the previous section, the results tend to argue in favor of a prolate tachocline thicker at high latitudes than at the pole, this may very well be due to the limits in resolution associated with the datasets used.

Time variations are also expected mainly because of the changes of the magnetic field strength and geometry during the solar cycle. Some evolution of the large scale circulation inside the layer or the action of internal gravity waves may also induce temporal changes or oscillations in the tachocline parameters. Small amplitude oscillations, with a period of 1.3 year, have recently been found in the rotation profile close to the boundaries of the tachocline (Howe et al. 2000). These results are based on both GONG and MDI data but have not been confirmed by independent analysis (Antia & Basu 2000). If these oscillations are real, the 1.3 year period remains unexplained (see Gough 2000). Because of the observed magnetic cycle, we rather expect the tachocline variations to be on the time scale of the solar cycle (11 years). A first attempt at detecting long term temporal variations in the tachocline have been made by Basu & Schou (2000) using 11 sets of MDI observations covering 72 days each between July 1996 and April 1998. They found no clear changes in the width or position of the tachocline but gave some hints that the position may move slightly outwards with increased activity.

5.2. The observations: LOWL instrument

The LOWL instrument is a Doppler imager based on a Potassium Magneto-Optical Filter that has been operating on Mauna Loa, Hawaï since 1994 (see Tomczyk et al. (1995) for a detailed description). The spectra were processed using the LOWL pipeline, which has been recently improved (Jiménez-Reyes 2001). Annual time series for degrees from $\ell = 0$ up to 99 have been created by using spherical harmonic masks and a Fast Fourier Transform

has been applied to each one of the time series. The average duty cycle over one year of observations was around 20%. A detailed description of the fitting method can be found in Jiménez-Reyes et al. (2001).

5.3. Method: inversion and deconvolution

In general, inferring the internal rotation rate from observed splittings requires inverting the integral Eq. 4. In order to achieve this, we have used a 2D Least-Squares (LS) method (Corbard et al. 1997). Because the method is linear, the solution at any point (r_0, ν_0) is linear combination of observations $< \Omega(r_0, \nu_0) >= \sum c_{nlm}(r_0, \mu_0)\delta_{nlm}$. From Eq. 4 this is equivalent to taking a linear combination of the kernels K_{nlm}. The resulting kernel

$$\kappa(r_0, \mu_0, r, \mu) = \sum c_{nlm}(r_0, \mu_0) K_{nlm}(r, \mu) \qquad (6)$$

is called averaging kernel because it follows from the previous equations that the quantity $< \Omega(r_0, \mu_0) >$ is an average of the rotation rate in a domain defined by $\kappa(r_0, \mu_0, r, \mu)$:

$$< \Omega(r_0, \mu_0) >= \int_0^1 \int_0^1 \kappa(r_0, \mu_0, r, \mu)\Omega(r, \mu) dr d\mu \quad (7)$$

The difficulties arise from the fact that Eq. 4 is an ill-posed problem with no unique solution. In this global LS method, we therefore need to introduce some a-priori knowledge on the rotation profile in order to *regularize* the solution and avoid strong oscillations. Nevertheless, this Regularized Least Squares (RLS) approach prevents us from recovering accurately sharp gradients as those expected in the tachocline because they do not conform to the global smoothness a-priori introduced. However we can then use our knowledge of the resolution in order to reach a better estimate of the tachocline width. The basic idea is to approximate Eq. 7 by a convolution equation (Charbonneau et al. 1999a; Corbard et al. 1998). If the the tachocline profile after inversion is approximated by an erf function of width w and if the radial part of averaging kernel is approximated by a Gaussian of standard deviation σ_r, then the 'true width' w_c can be obtained by:

$$w_c = \sqrt{w^2 - 8\sigma_r^2} \qquad (8)$$

5.4. Results and Discussion

The results are shown in Fig. 4. The radial resolution obtained is about $0.12R_\odot$ at 60° and $0.10R_\odot$ at the equator. The angular resolution achieved with LOWL data is about 20° at the equator. Because the radial shear doesn't exist at about 30° and because the angular resolution decrease at high latitudes, we limit our analysis to two latitudes: the equator and 60°. Therefore, in the following, prolateness refers to the difference between the central position of the layer at these two latitudes.

Because the inferred width at the equator is always lower than the resolution we cannot use our simple model

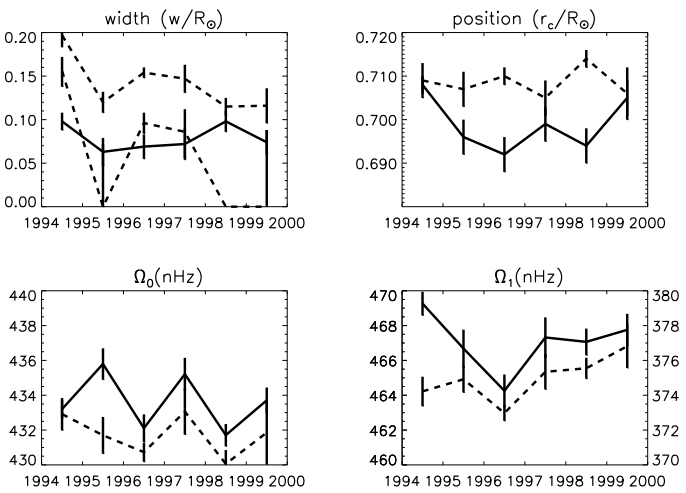

Figure 4. Inferred tachocline parameters as a function of time from LOWL observations. The full and dashed lines correspond respectively to fit at the equator and at 60°. *For the width (upper left), the lower dashed line is obtained after deconvolution using Eq. 8. The scale on the right of the lower right panel corresponds to the fit at* 60° *(dashed line).*

for deconvolving. This indicates that the width of the tachocline at the equator is probably lower than the local spacing of the grid used for the inversion i.e. $0.02R_\odot$. The same happens at 60° for 1995, 1998 and 1999. There is therefore no strong evidence of a systematically thicker tachocline at 60°. Moreover the errors reported on the plot are formal errors as obtained from the fits but Monte Carlo simulations have been carried out which suggest that the uncertainties on the inferred widths after the whole inversion process are between ± 0.02 and $\pm 0.03R_\odot$.

The center of the tachocline is always found deeper at the equator than at 60°. The variation of this prolateness with time is shown in Fig. 5. The maximum of prolateness found is about $0.02R_\odot$ and no prolateness is found the first and last years. These differences and their fluctuations are nevertheless very small and are also very sensitive to the inversion parameters chosen and especially the amount of regularization used. The lower panel of Fig. 5 illustrates this point by showing that, with a more regularized inversion, the maximum prolateness is about $0.03R_\odot$ and a minimum is no longueur found for 1994.

Generally speaking, we do not find from this analysis evidence of any general trend or significant oscillation in the tachocline parameters during the ascending phase of the actual solar cycle. In particular we do not find an outward trend for the central position as suggested by Basu & Schou (2000). Nevertheless, a prolateness of the layer is observed every year which allow us to, at least, exclude an oblate tachocline. The amount of prolateness found between the equator and 60° is however of the same order of magnitude than the uncertainties on the width of the layer. We can therefore only set an upper limit for the prolateness which is around $0.03R_\odot$. This is in good

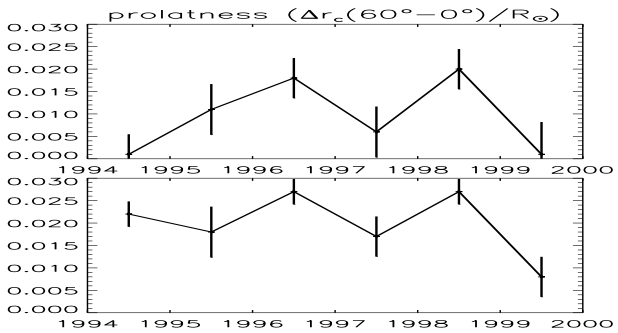

Figure 5. Difference between the central position inferred at 60° *and at the equator, as a function of time. The upper plot corresponds to the results of Fig. 4 while the lower plot corresponds to inversion using stronger regularization.*

agreement with previous estimates of Antia et al. (1998) ($0.004<\Delta_{rc}/R_\odot<0.023$) and Charbonneau et al. (1999a) ($\Delta_{rc} \simeq 0.023 R_\odot$). Using the shallow-water model and shape curves as shown in Fig. 3, this would correspond to a toroidal magnetic field strength of about 0.1MG if it is located in the overshoot layer or about 0.6MG if it is located in the radiative interior. If the toroidal field is concentrated in bands migrating towards the equator during the ascending phase of the cycle, one would expect, from the shallow water model, a decreasing prolateness for the period of LOWL observations. This general trend cannot be excluded but is not observed from our preliminary analysis of LOWL data.

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REFERENCES

Antia H.M., Basu S., 2000, ApJ, 541, 442

Antia H.M., Basu S., Chitre S.M., 1998, MNRAS, 298, 543

Basu S., 1997, MNRAS, 288, 572

Basu S., Antia H.M., 1997, MNRAS, 287, 189

Basu S., Schou J., 2000, Solar Phys., 192, 481

Boruta N., 1996, ApJ, 458, 832

Brun A.S., Turck-Chièze S., Zahn J.P., 1999, ApJ, 525, 1032

Canuto V.M., 1997, ApJ, 489, L71

Chaboyer B., Demarque P., Pinsonneault M.H., 1995, ApJ, 441, 865

Charbonneau P., MacGregor K.B., 1993, ApJ, 417, 762

Charbonneau P., Christensen-Dalsgaard J., Henning R., et al., 1999a, ApJ, 527, 445

Charbonneau P., Dikpati M., Gilman P.A., 1999b, ApJ, 526, 523

Christensen-Dalsgaard J., Monteiro M.J.P.F.G., Thompson M.J., 1995, MNRAS, 276, 283

Corbard T., Berthomieu G., Morel P., et al., 1997, A&A, 324, 298

Corbard T., Berthomieu G., Provost J., Morel P., 1998, A&A, 330, 1149

Dikpati M., Gilman P.A., 1999, ApJ, 512, 417

Dikpati M., Gilman P.A., 2000a, ApJ, in press

Dikpati M., Gilman P.A., 2000b, ApJ, submitted

Elliott J.R., 1997, A&A, 327, 1222

Ferriz-Mas A., 1996, ApJ, 458, 802

Ferriz-Mas A., Schmitt D., Schuessler M., 1994, A&A, 289, 949

Garaud P., 1999, MNRAS, 304, 583

Gilman P.A., 2000, Solar Phys., 192, 27

Gilman P.A., Fox P.A., 1997, ApJ, 484, 439

Gilman P.A., Fox P.A., 1999, ApJ, 510, 1018

Gough D., 2000, science, 287, 2434

Gough D.O., McIntyre M.E., 1998, Nat, 394, 755

Howe R., Christensen-Dalsgaard J., Hill F., et al., 2000, science, 287, 2456

Hujeirat A., Yorke H.W., 1998, New Astronomy, 3, 671

Jiménez-Reyes S.J., 2001, in progress, Ph.D. thesis, La Laguna University

Jiménez-Reyes S.J., Corbard T., Palle P.L., Tomczyk S., 2001, In: SOHO 10/ GONG 2000 Workshop, Helio- and Asteroseismology at the Dawn of the Millennium, vol. SP-464, ESA Publications Division, Noordwijk, The Netherlands

Kim E.J., MacGregor K.B., 2000, ApJ, submitted

Kosovichev A.G., 1996, ApJ, 469, L61

Kumar P., Talon S., Zahn J., 1999, ApJ, 520, 859

MacGregor K.B., Charbonneau P., 1999, ApJ, 519, 911

Mestel L., Weiss N.O., 1987, MNRAS, 226, 123

Monteiro M.J.P.F.G., Christensen-Dalsgaard J., Thompson M.J., 1998, In: Korzennik S.G., Wilson A. (eds.) Structure and Dynamics of the Interior of the Sun and Sun-like Stars, vol. SP- 418, 495–498, ESA Publications Division, Noordwijk, The Netherlands

Petrovay K., 2001, this proceeding

Richard O., Vauclair S., Charbonnel C., Dziembowski W.A., 1996, A&A, 312, 1000

Rudiger G., Kitchatinov L.L., 1997, Astronomische Nachrichten, 318, 273

Schatzman E., Zahn J.P., Morel P., 2000, A&A, submitted

Schüssler M., Caligari P., Ferriz-Mas A., Moreno-Insertis F., 1994, A&A, 281, L69

Spiegel E.A., Zahn J.P., 1992, A&A, 265, 106

Tomczyk S., Streander K., Card G., et al., 1995, Solar Phys., 159, 1

Zahn J.P., 1991, A&A, 252, 179